

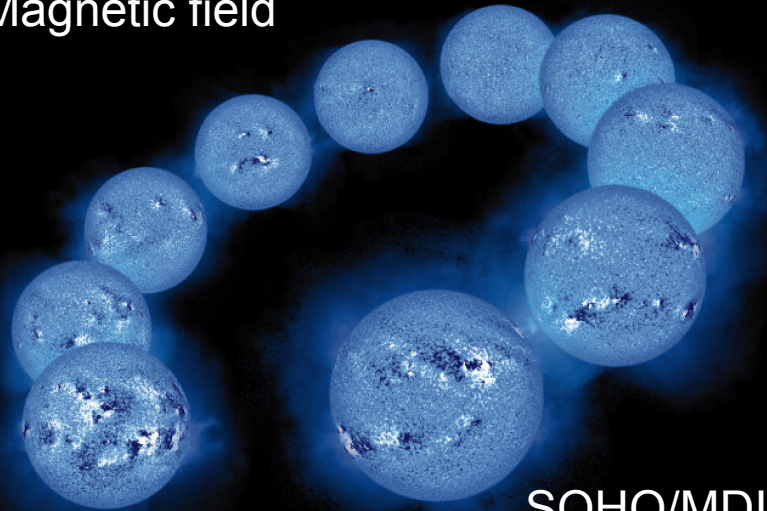
Solar wind acceleration

Shinsuke Imada

National Astronomical Observatory Japan

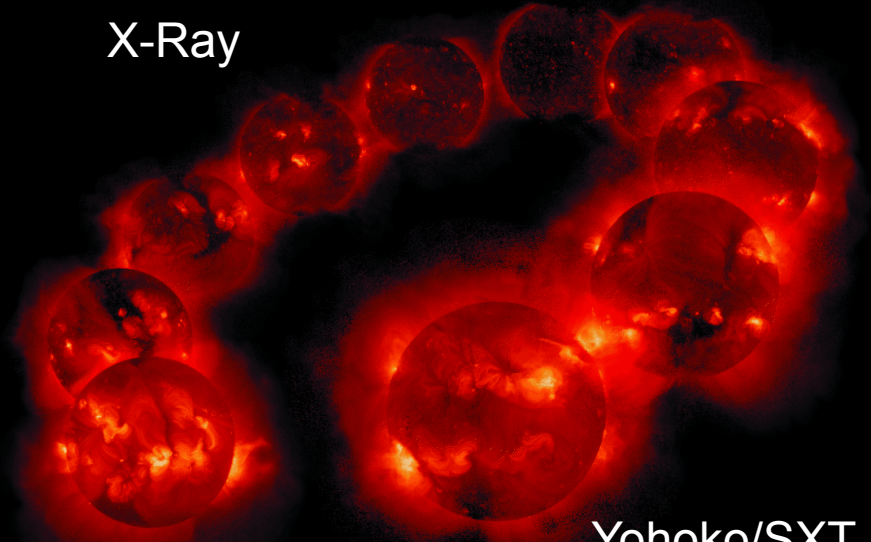
Remote-sensing Observation before Hinode

Magnetic field



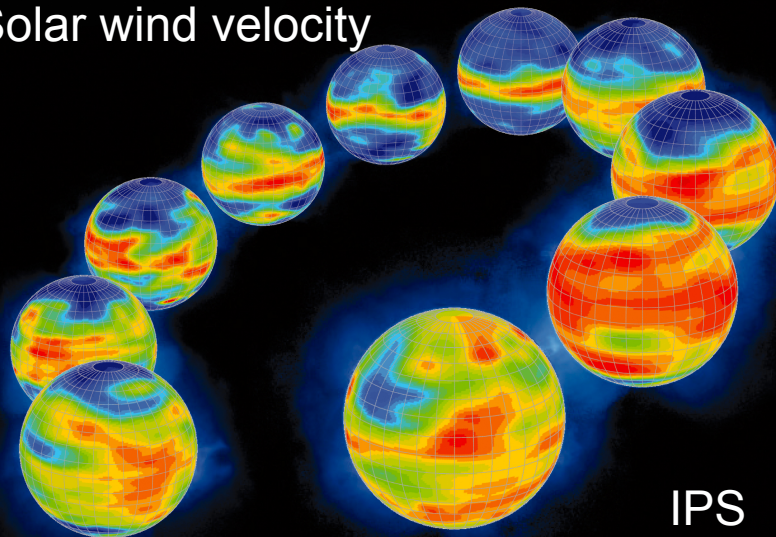
SOHO/MDI

X-Ray



Yohoko/SXT

Solar wind velocity

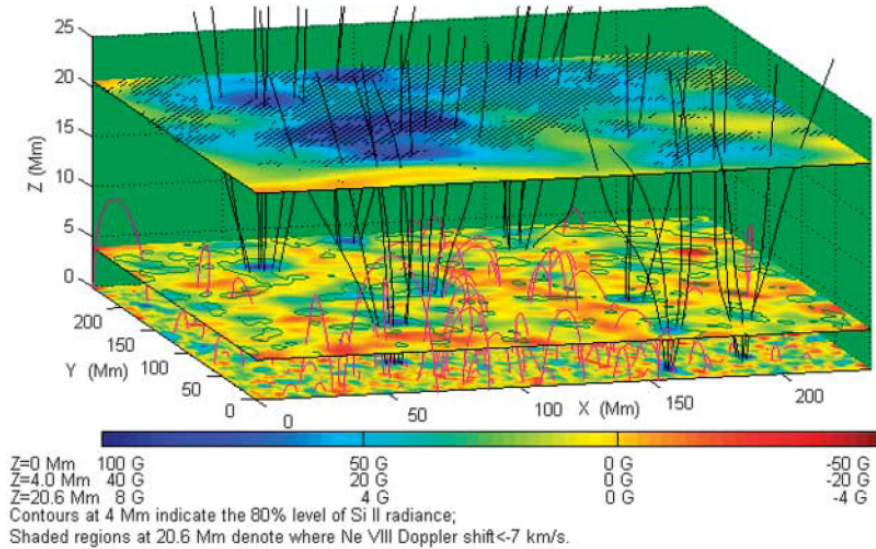


IPS

Solar wind structure dramatically changes with solar activity.

At the solar minimum phase, the fast wind is concentrate on the polar region. On the other hand, at the maximum phase it is not so much simple.

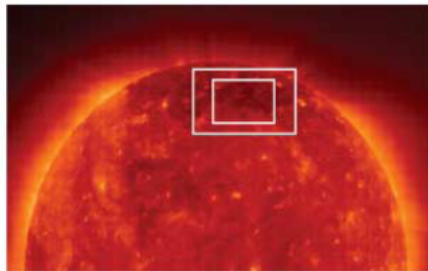
Polar Coronal hole (fast solar wind)



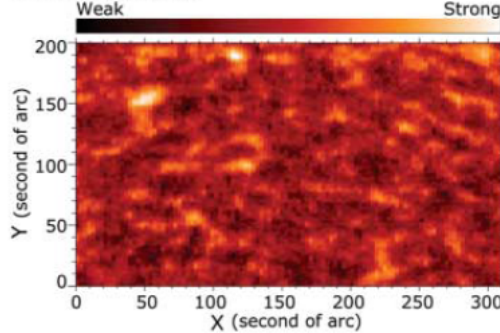
Soho/SUMER observation

Tu et al., 2005

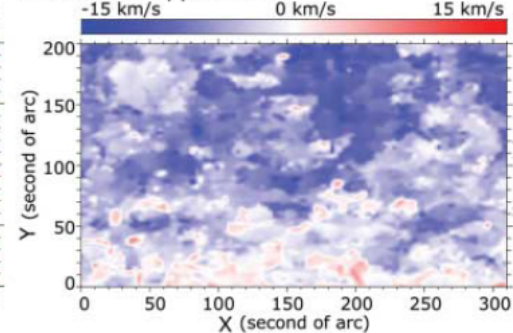
A EIT 19.5 nm



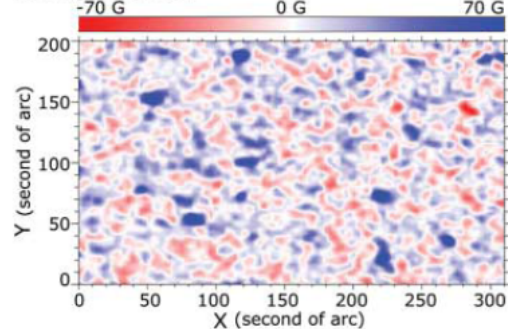
C Si II Radiance



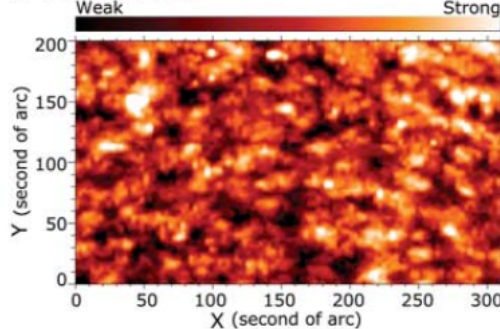
E Ne VIII Doppler shift



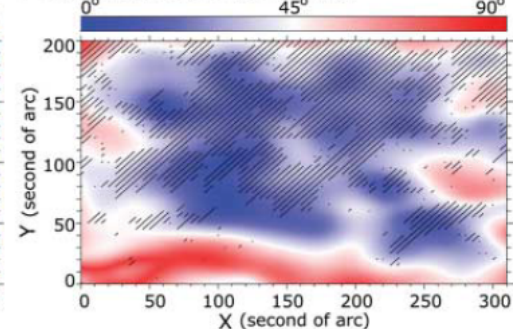
B Magnetic field



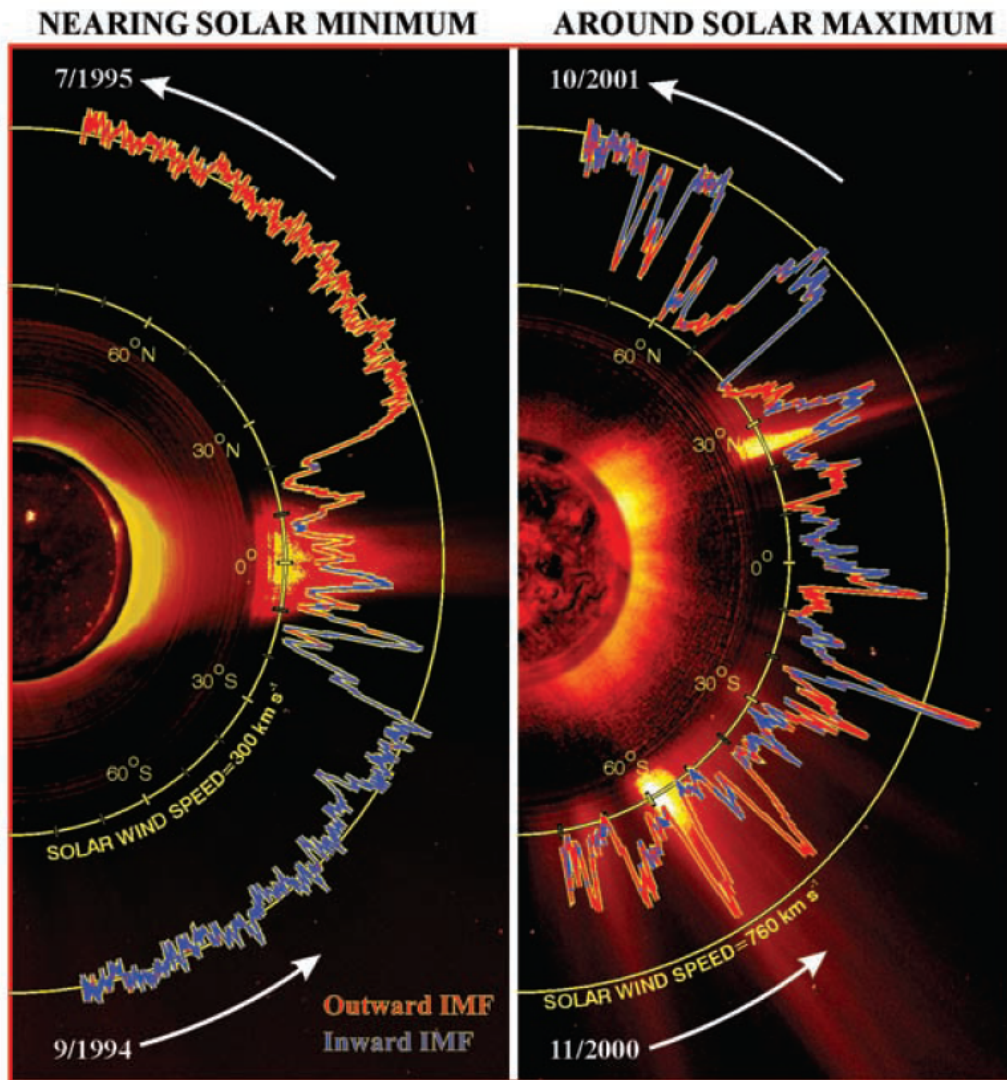
D C IV Radiance



F Field inclination at 20.6 Mm



In-situ measurement



Radial distance: solar wind velocity

Red: Outward IMF

Blue: Inward IMF

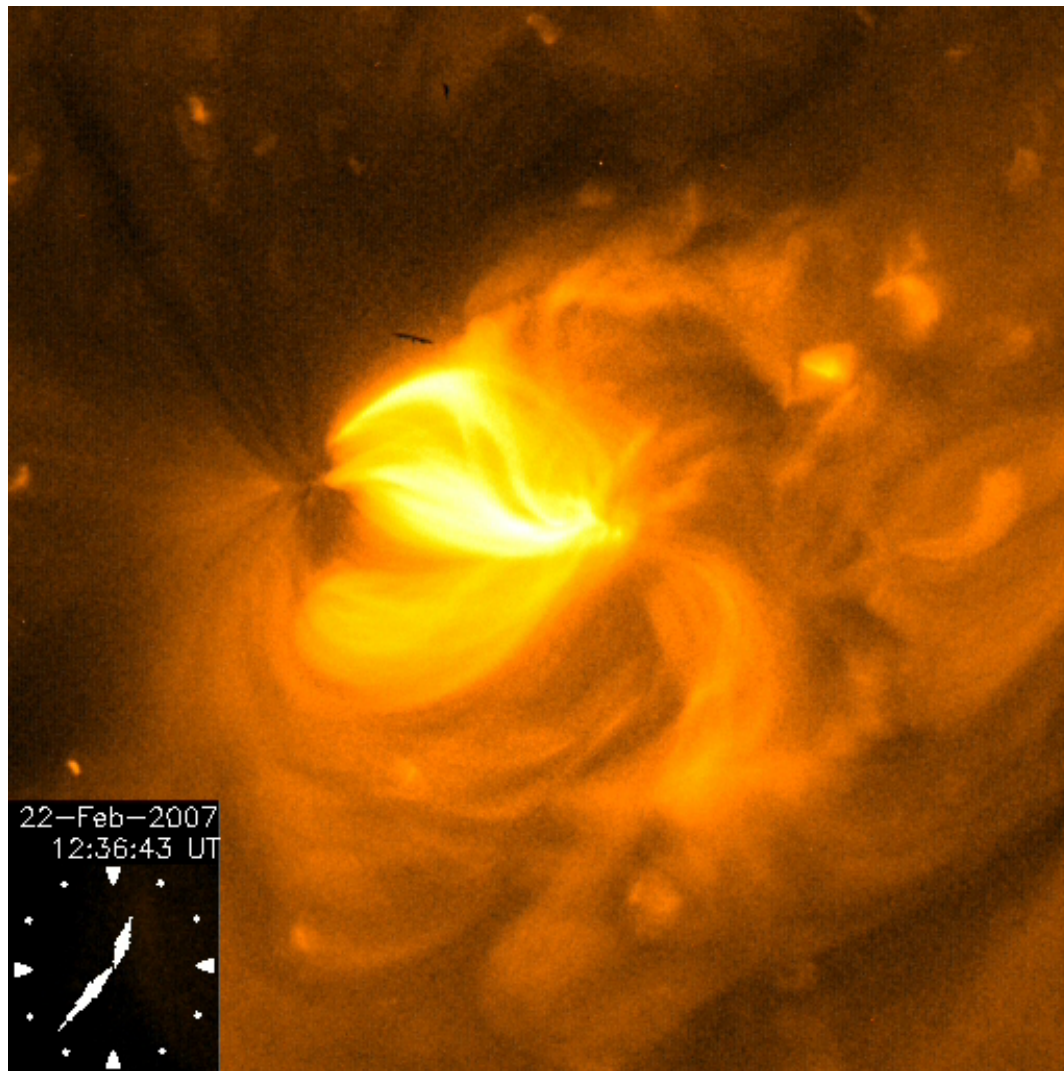
Fast solar wind show the steady flow ~700km/sec, magnetic fields direction is also one polarity

On the other hand, slow wind show the fluctuate structure.

Not only the velocity but also magnetic field direction.

Ulysses observation
McComas et al., 2000

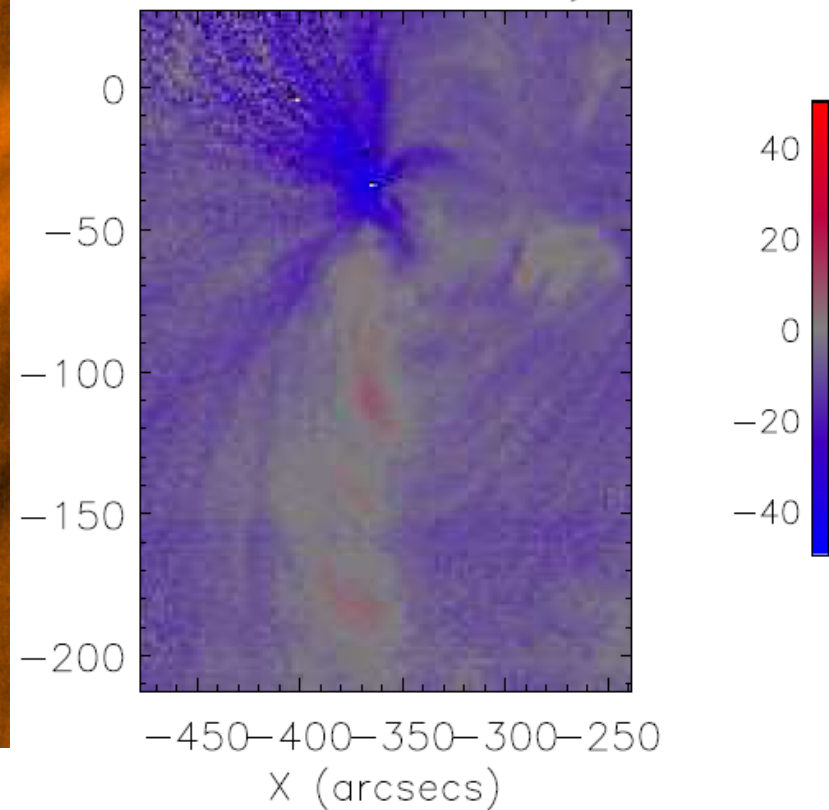
Hinode Observation (Slow Wind)



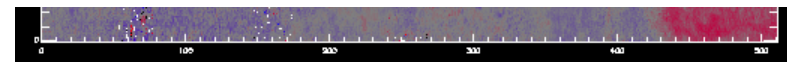
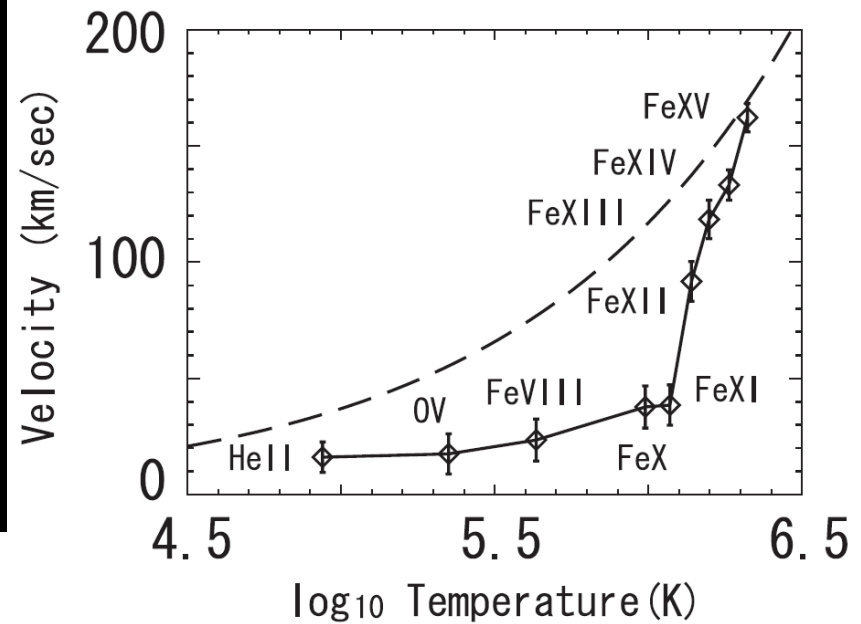
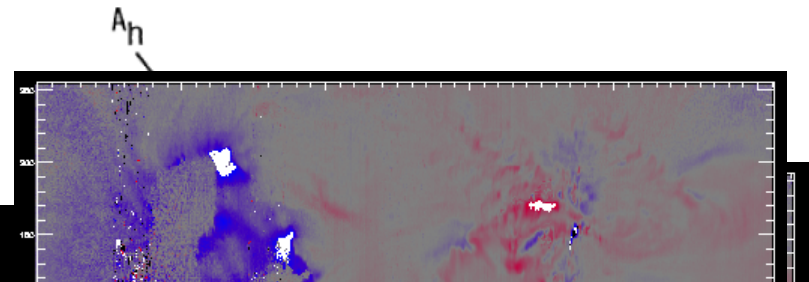
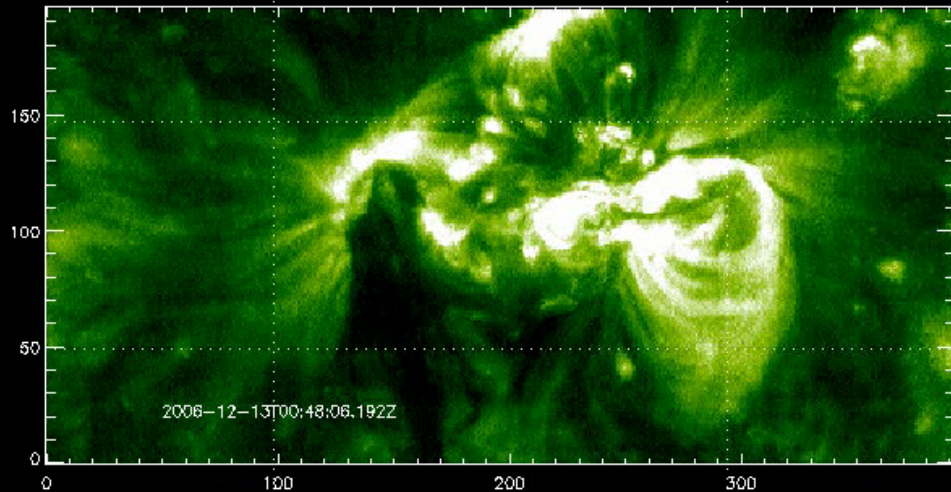
Sakao et al, 2007

Harra et al, 2008

EIS 195, velocity



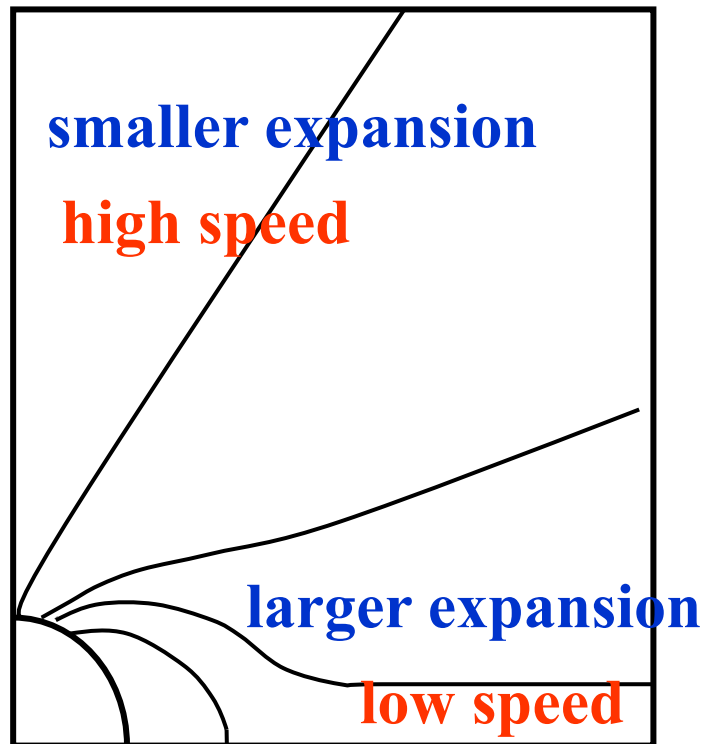
Transient Coronal Hole (Dimming)



Imada et al, 2007

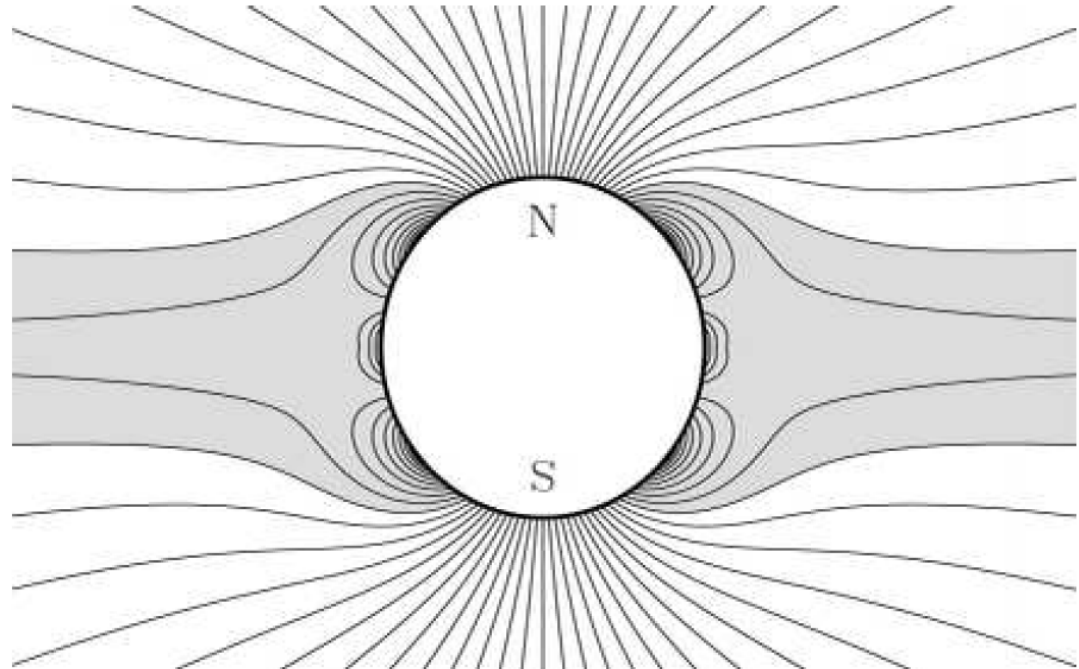
What is the difference between Slow and Fast?

Magnetic field configuration



RADIAL DISTANCE

Wang and Sheeley 1990



Larger expansion \rightarrow Heating near the Sun

\rightarrow Temperature gradient is large

\rightarrow Large density (not mass flux) into interplanetary

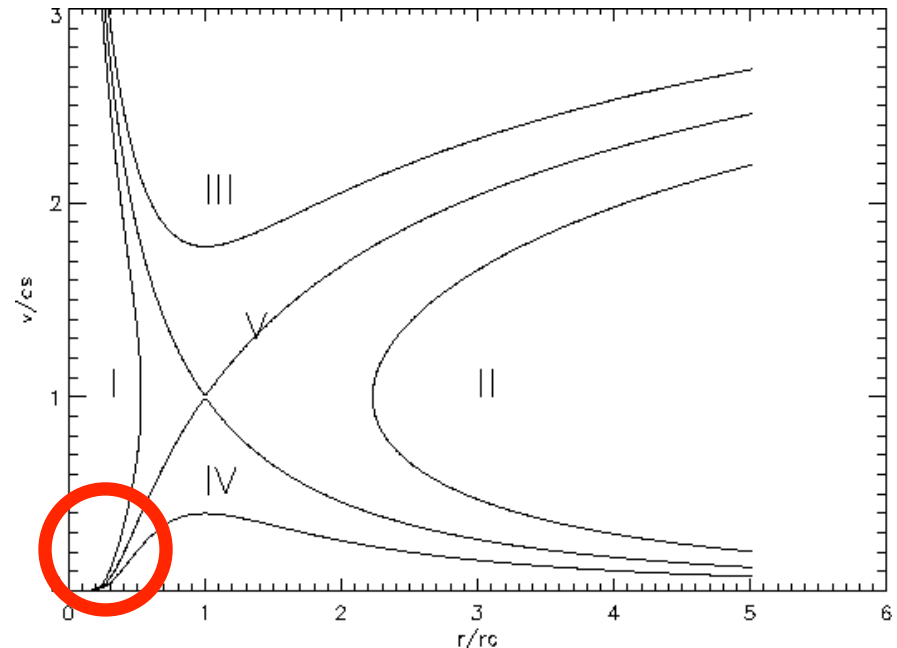
This is one of the reason why we can observe slow wind region by Hinode

What is for Solar-C?

- The origin of steady fast solar wind.
(injection to solar wind)

This is very difficult because the density is very low. Need high-throughput EUV spectrograph (This will be discussed today).

Parker's solar wind model



Why fast solar wind is so steady?

- Slow solar wind origin is not steady
→ This will be understood in more detail with Hinode Observation.
- Is fast solar wind origin steady?
→ maybe not
 - Magnetic field in the Polar region (Tsuneta-san's Talk)
 - X-ray jets & Plume (Cirtain-san's Talk)
 - Chromosphere (Tomorrow Talk)

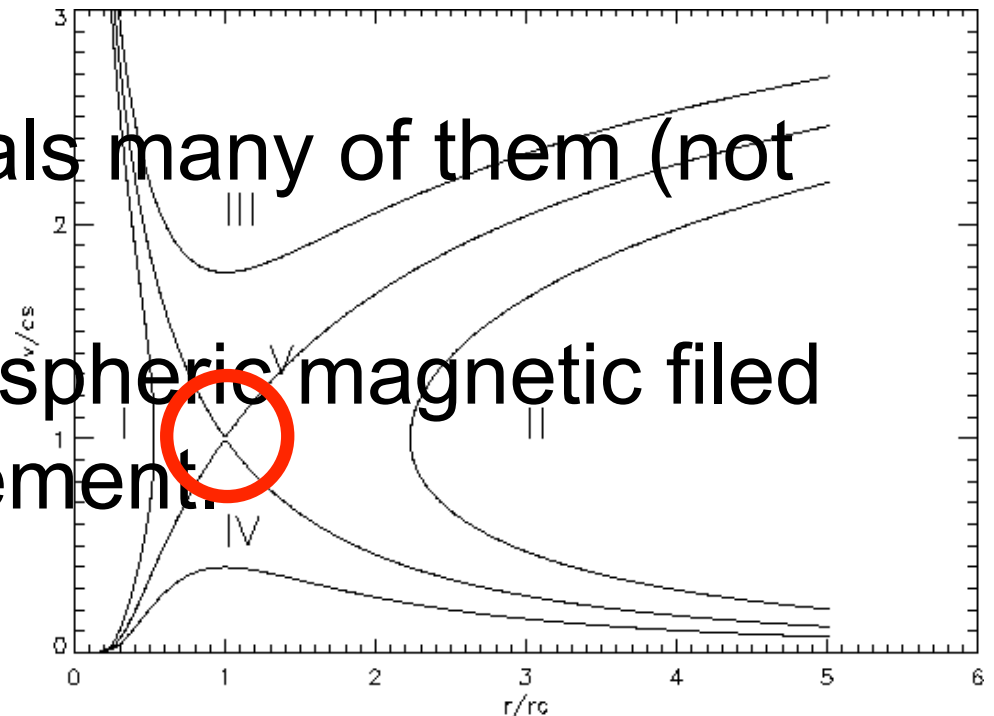
What is for Solar-C?

- Can we observe critical point (velocity \sim sound speed)?
→ in-situ measurement helpful

It is up to payload..

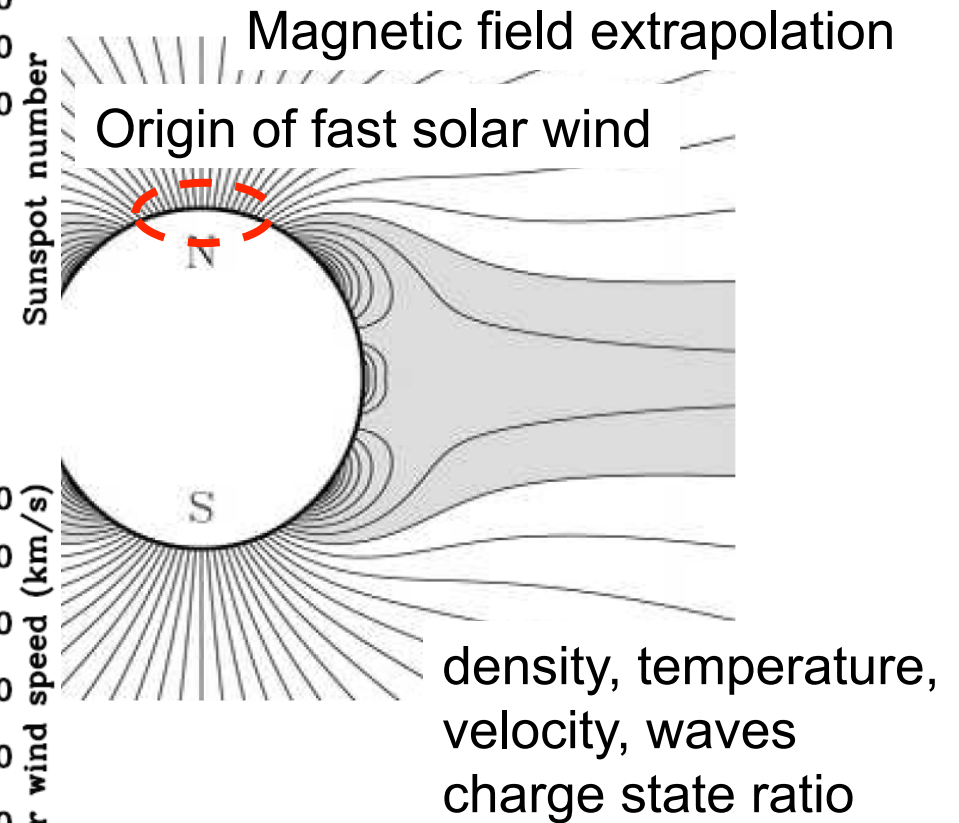
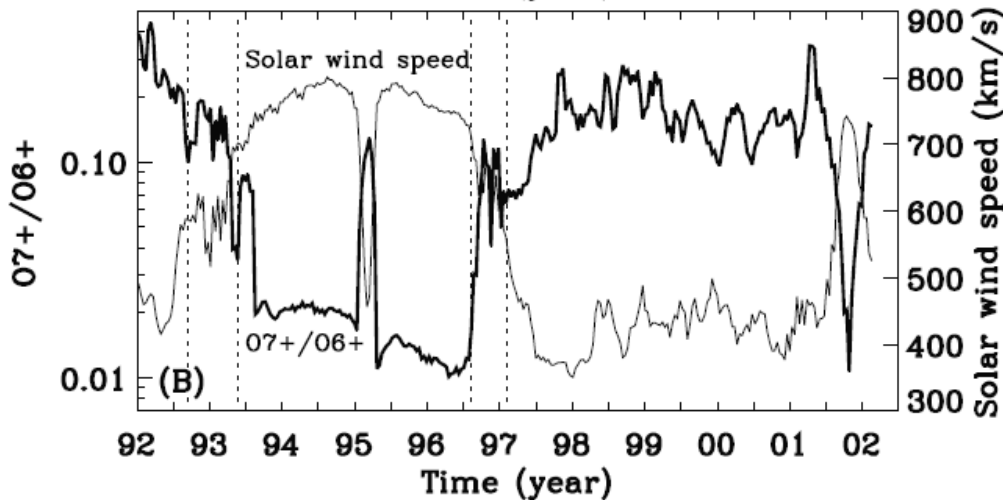
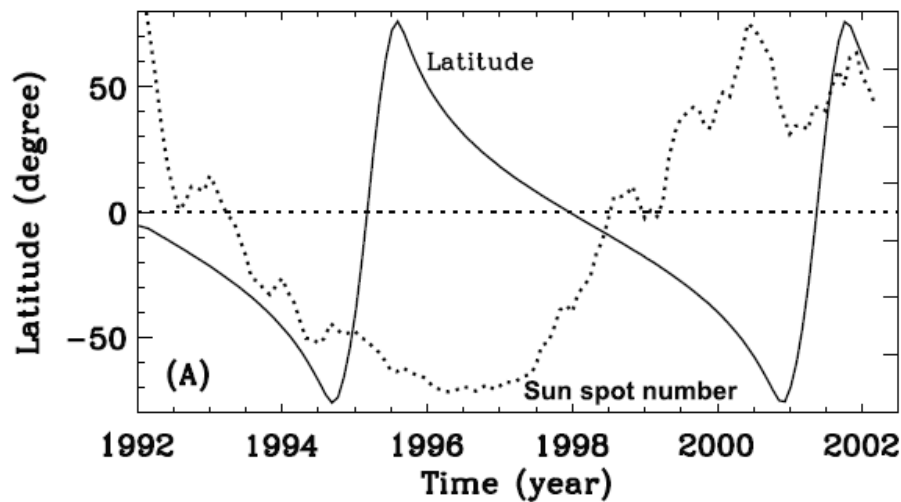
Actully, Ulysses reveals many of them (not statistically).

Combination of photospheric magnetic filed and in-situ measurement.



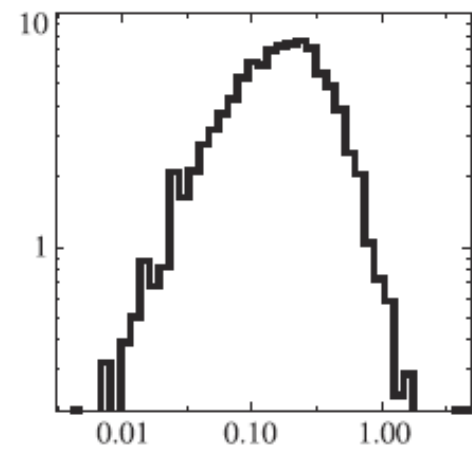
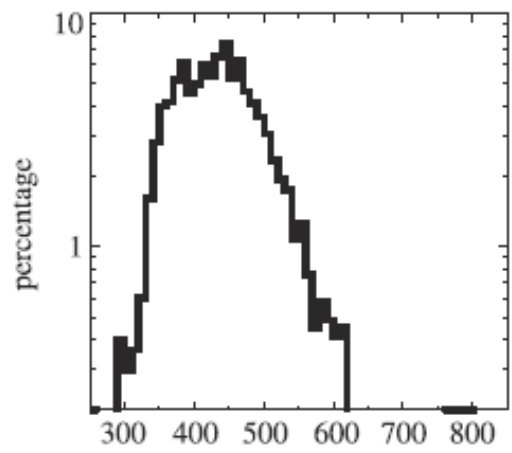
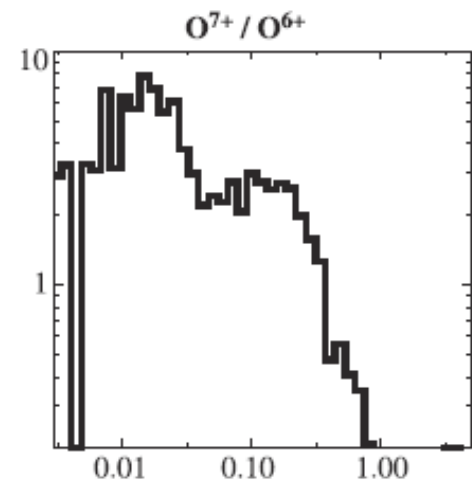
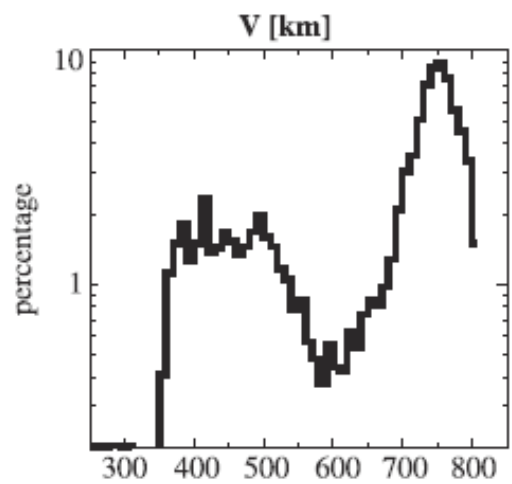
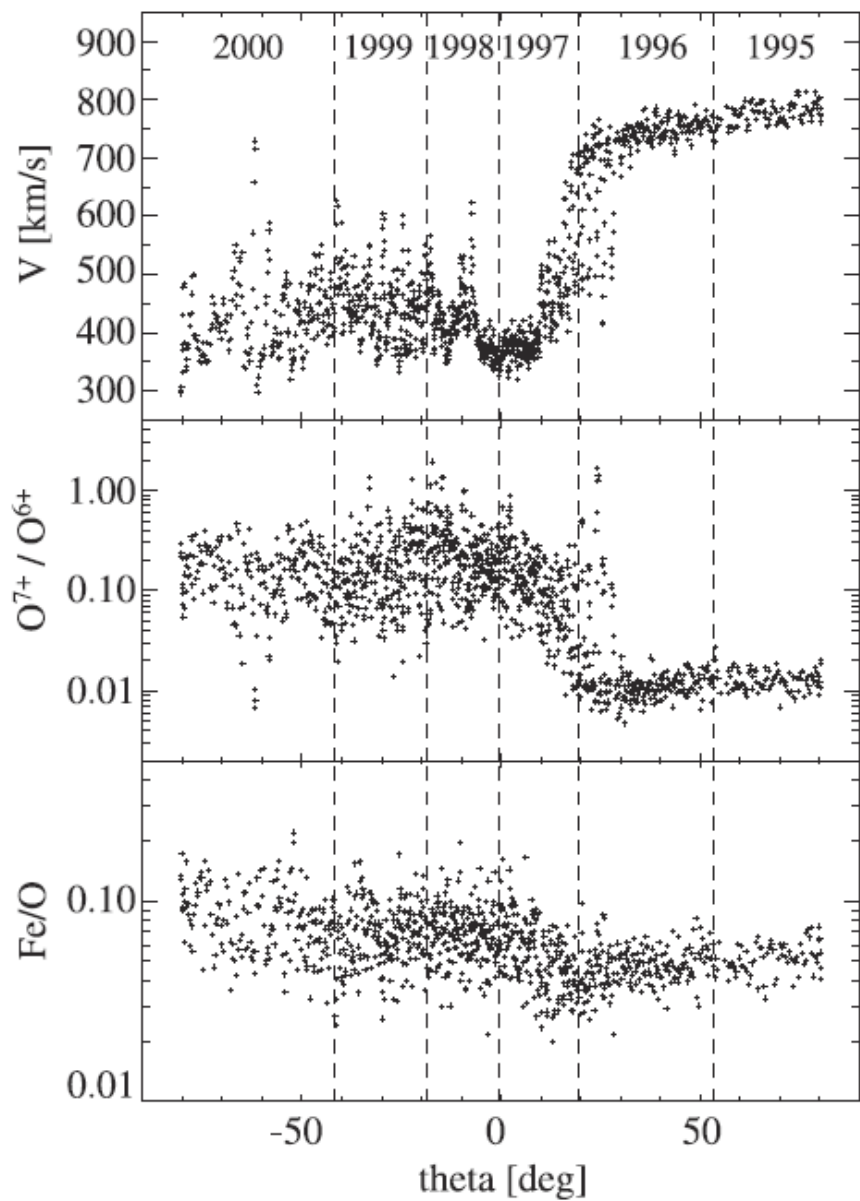
Combination remote-sensing and in-situ

Zhang et al, 2003

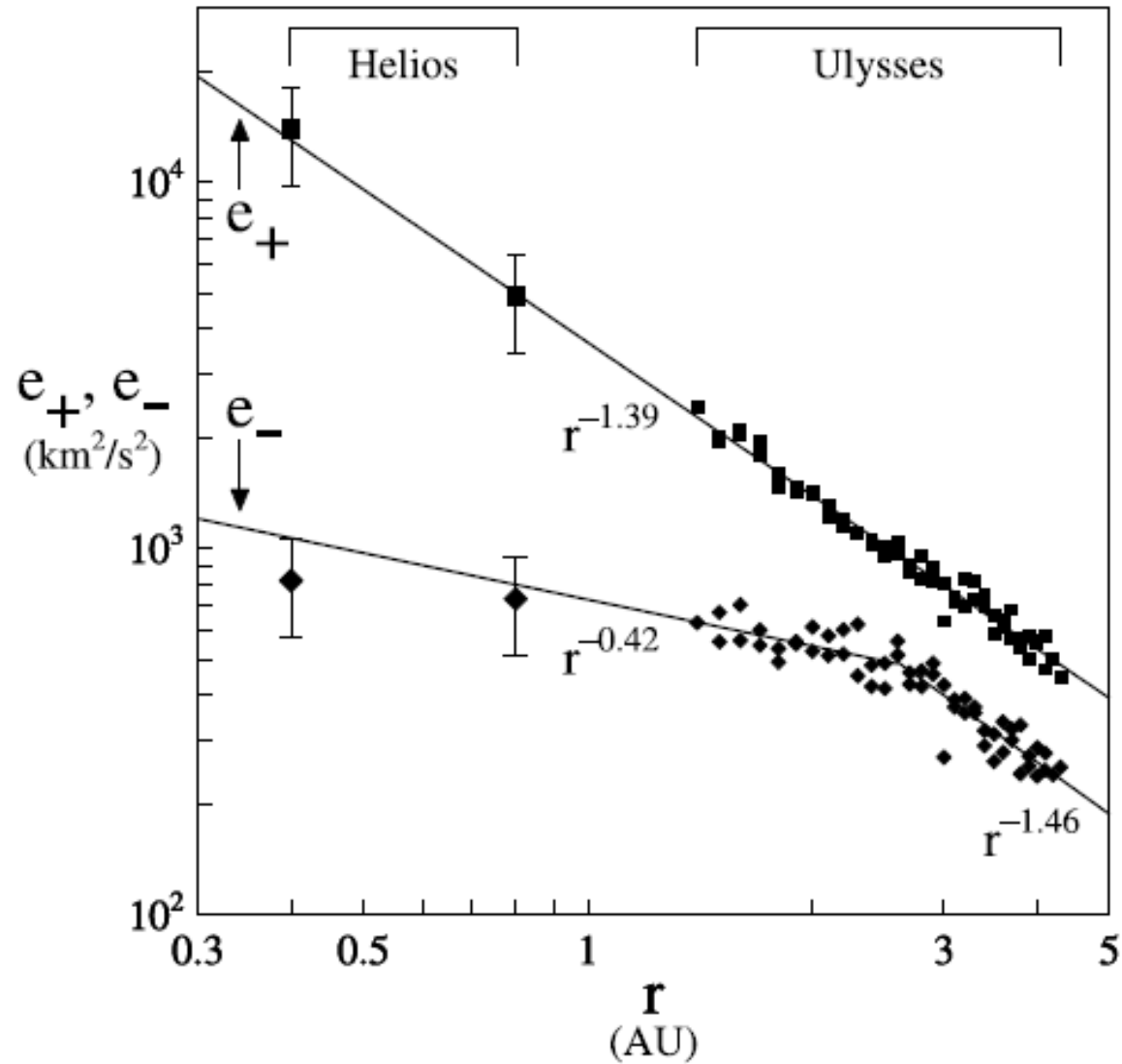


Summary

- Origin of fast solar wind. “steady or unsteady?”
- We need high throughput EUV spectrometer to observe darker region
- Combination with in-situ measurement is helpful for understanding the critical point.
It is up to payload.
- This analysis need precise modeling.



Inward and Outward Alfenwave

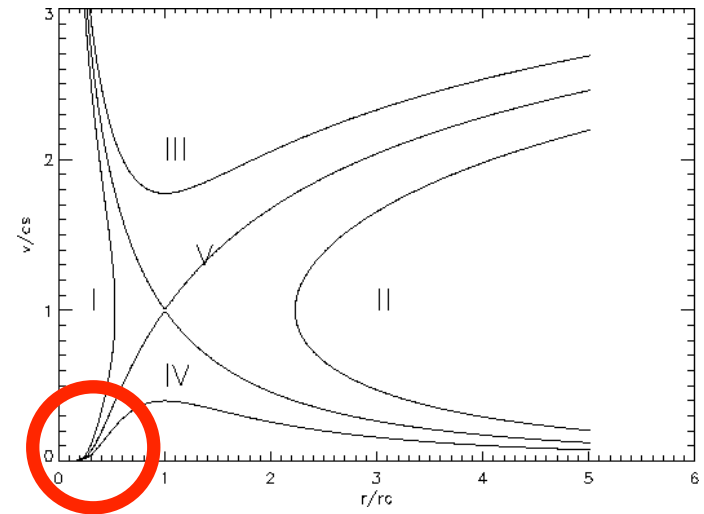


Parker's Solar wind model

$$4\pi mnur^2 = \text{constant}$$

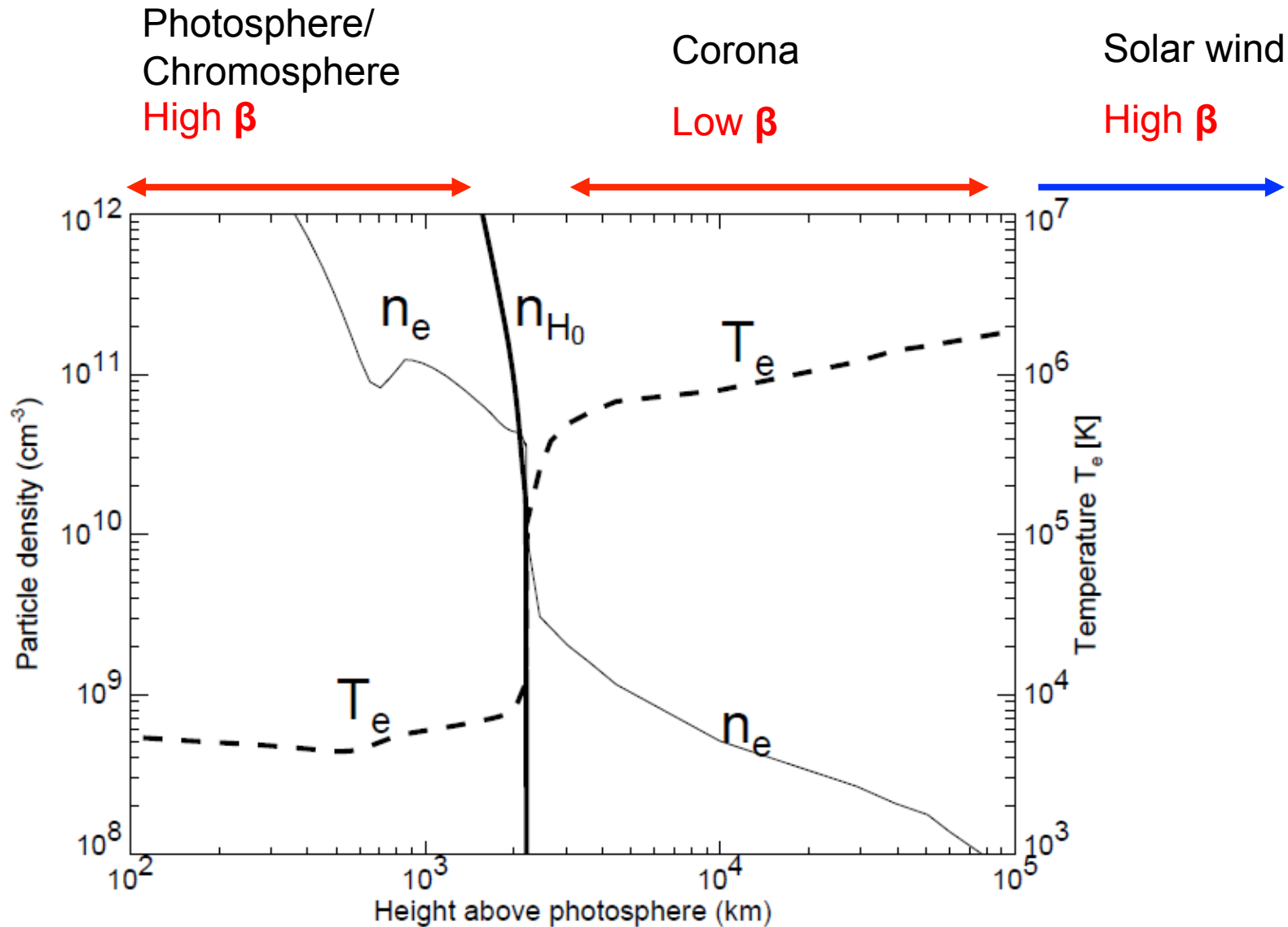
$$u \frac{du}{dr} + \frac{1}{mn} \frac{dp}{dr} + \frac{GM_S}{r^2} = 0$$

$$p = 2nkT$$

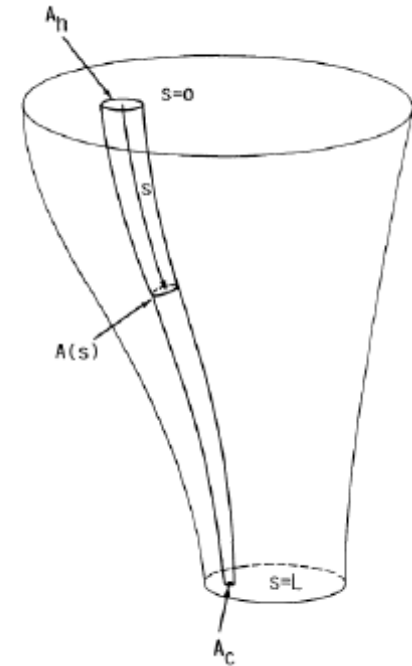
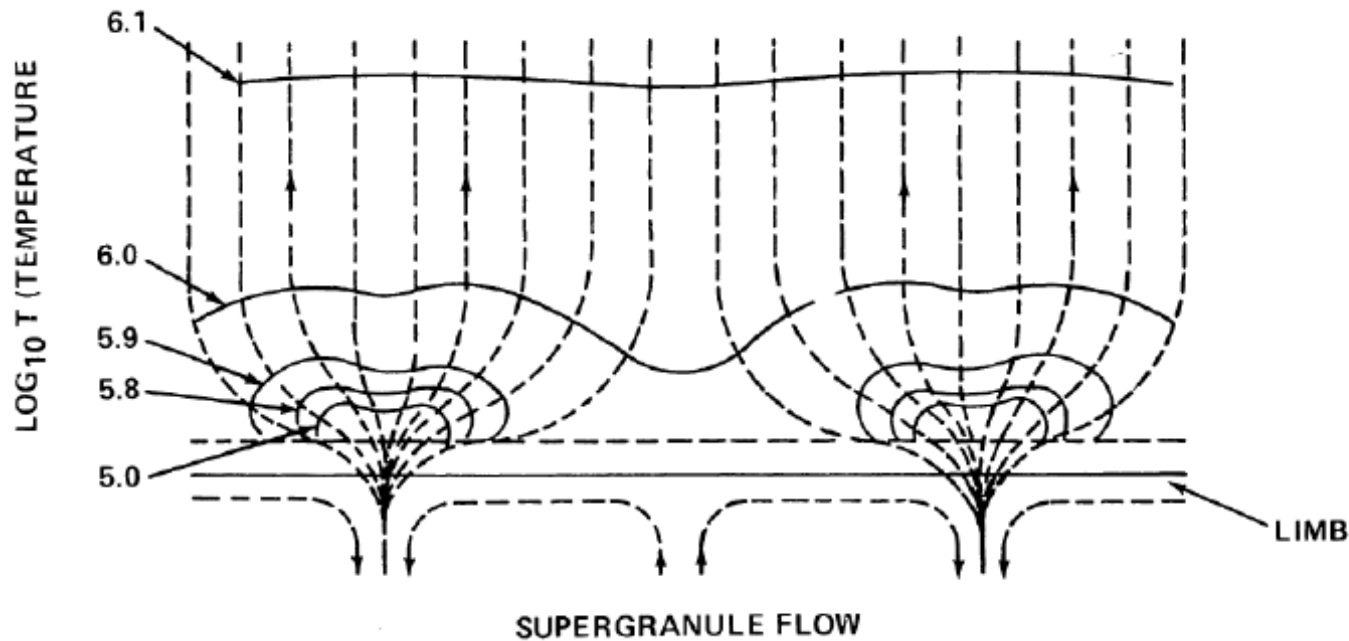


$$\frac{1}{u} \frac{du}{dr} = \left(\frac{2a^2}{r} - \frac{GM_S}{r^2} \right) / (u^2 - a^2)$$

From Photosphere to Corona

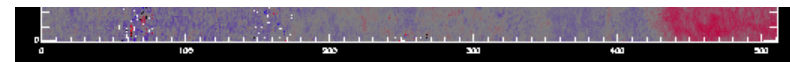
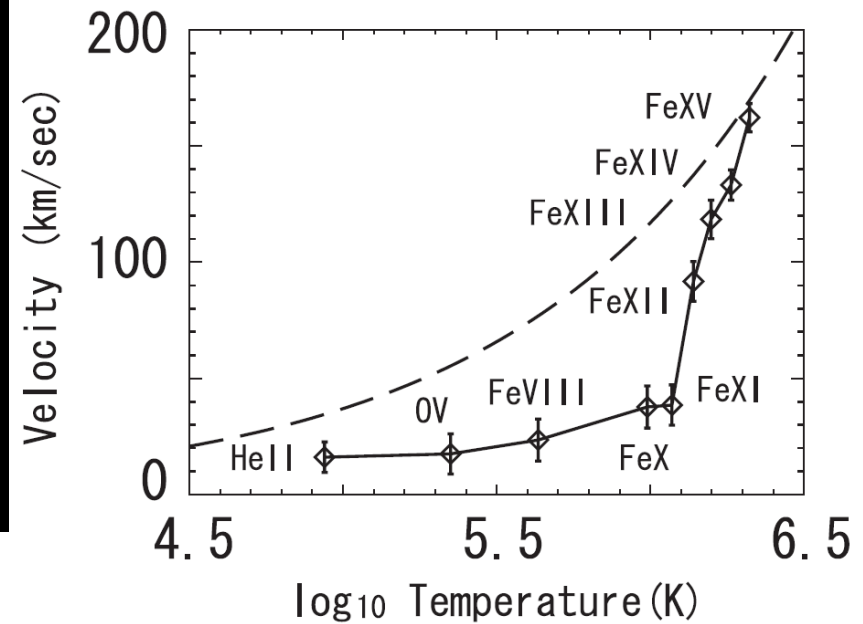
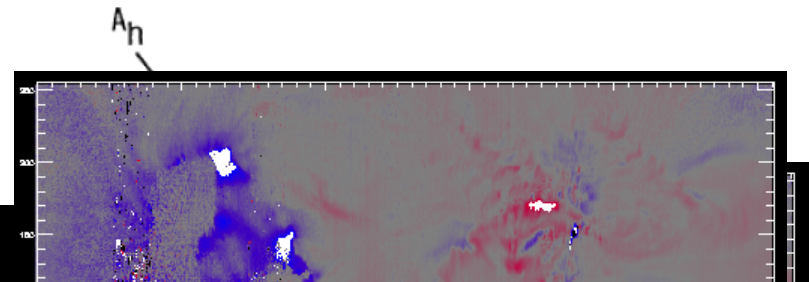
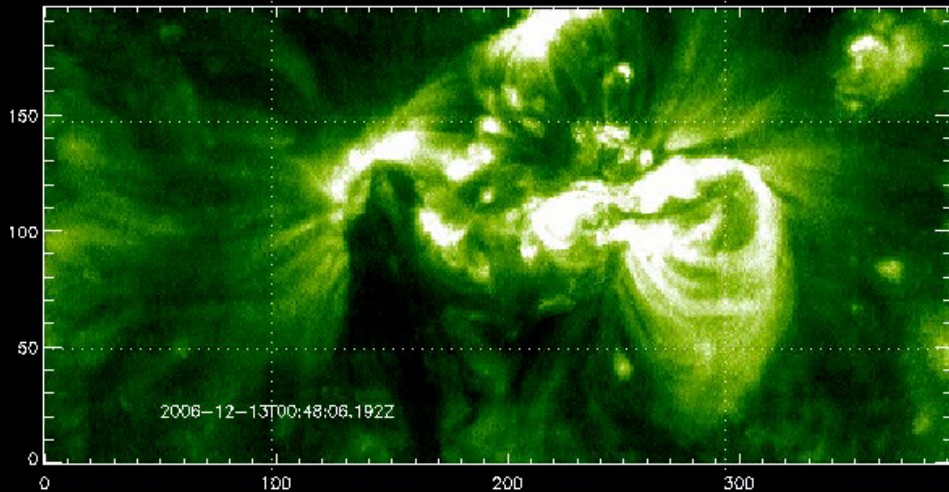


Coronal Funnel

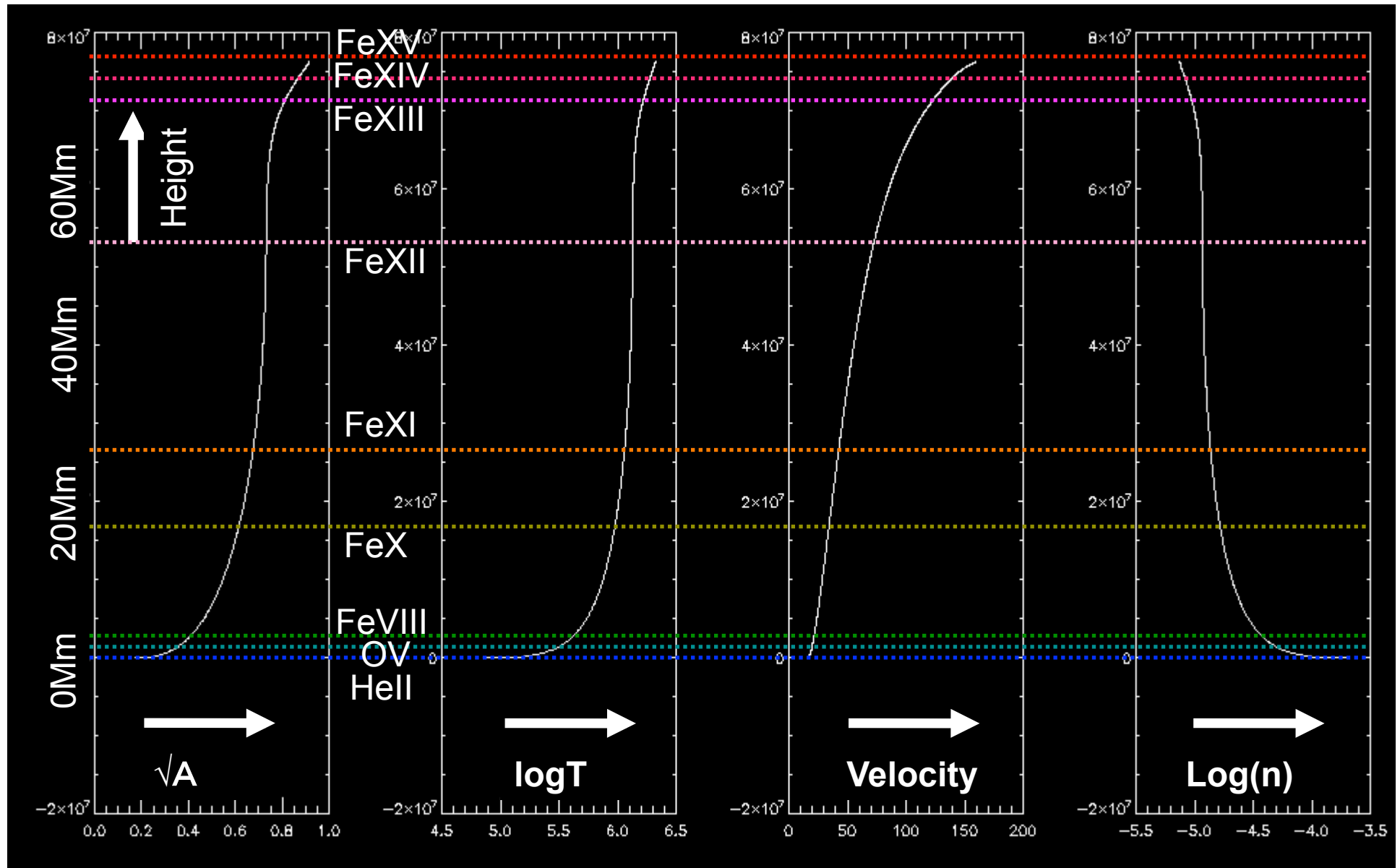


$$\left(v - \frac{c_s^2}{v}\right) \frac{\partial v}{\partial s} = -\frac{2k_B}{m} \frac{\partial T}{\partial s} + \frac{c_s^2}{A} \frac{\partial A}{\partial s} - g$$

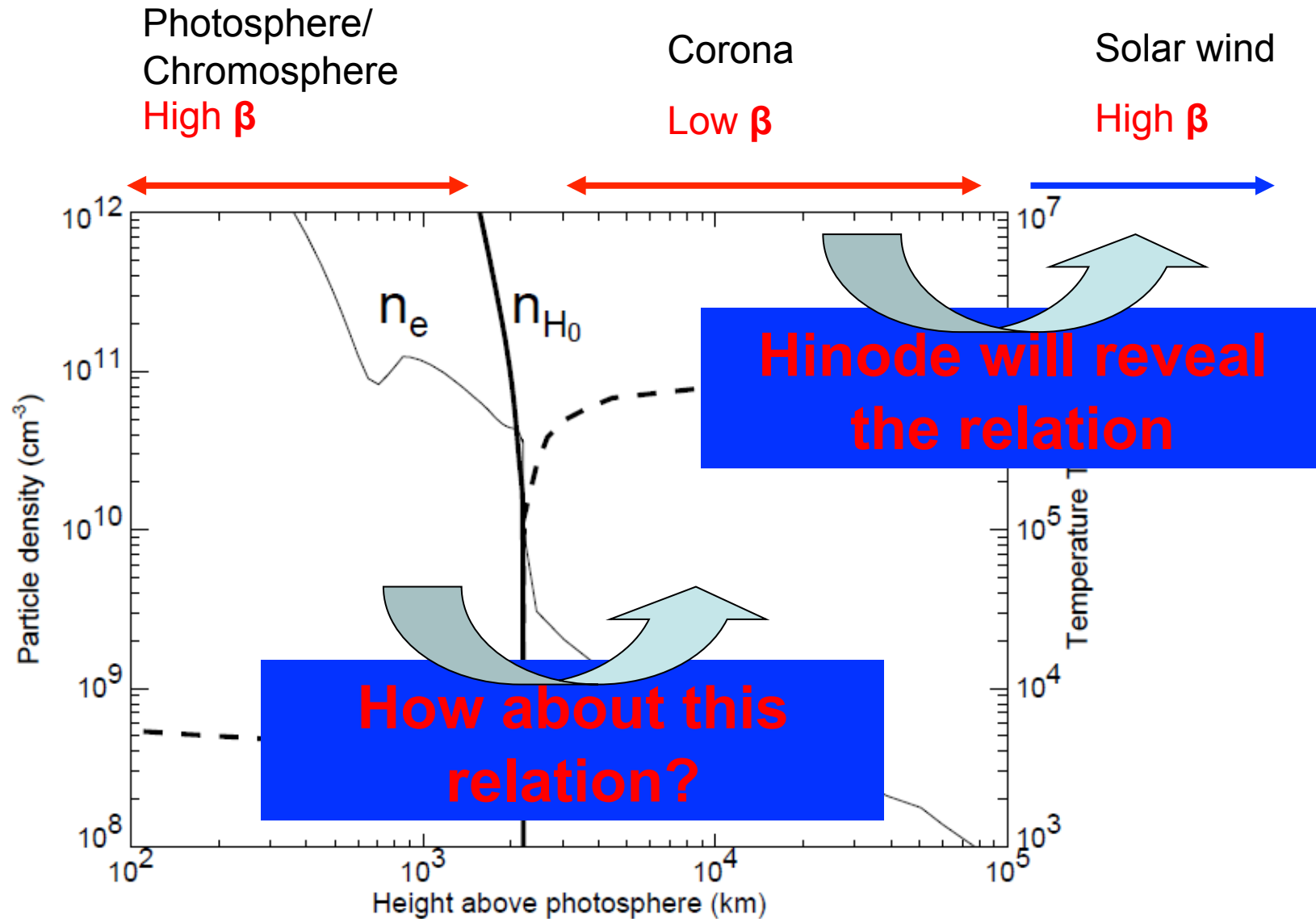
Transient Coronal Hole (Dimming)



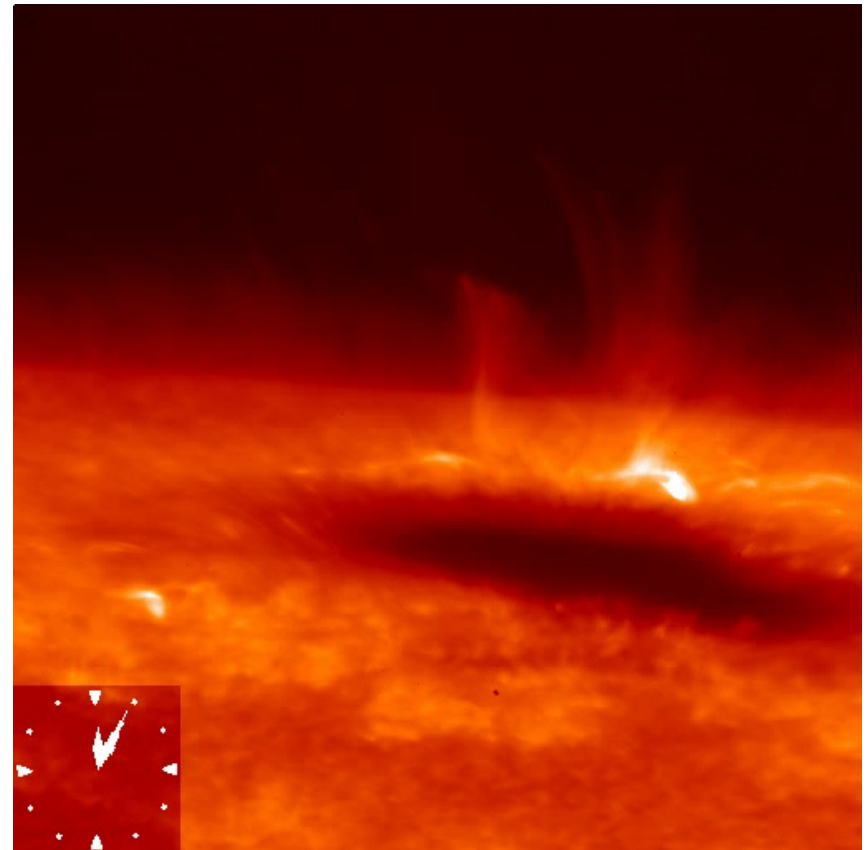
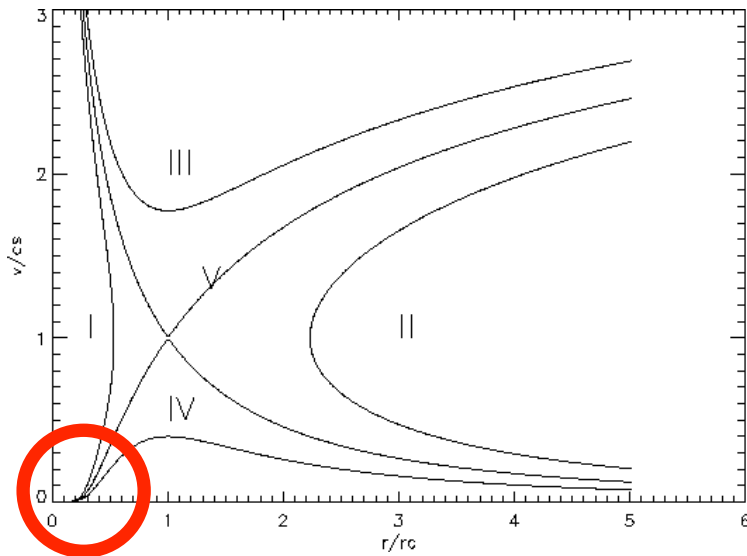
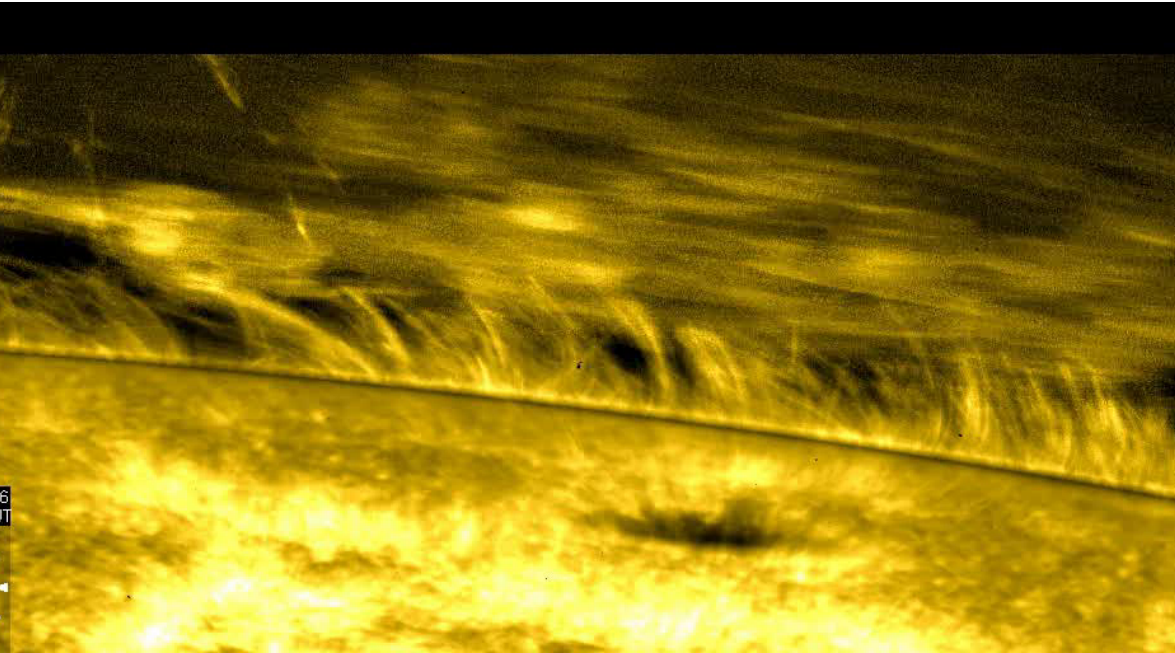
Extrapolation for Z-direction



From Photosphere to Corona



Steady or Unsteady injection?



For observing injection region

- High sensitivity
 - to observe very darker region
- Magnetic field measurement in chromosphere/transition region
 - to observe expansion factor
- High cadence observation
 - to observe dynamic features

Observe from Polar region is better..